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Signatures of quiet Sun reconnection events in Ca II, H α and Fe I.

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ABSTRACT

We use observations of quiet Sun (QS) regions in the H α 6563 Å, Ca II 8542 Å and Fe I 6302 Å lines. We observe brightenings in the wings of the H α and Ca II **combined with observations of the interacting magnetic concentrations observed in the Stokes signals of Fe I**. These brightenings are similar to Ellerman bombs (EBs), i.e. impulsive bursts in the wings of the Balmer lines which leave the line cores unaffected. Such enhancements suggest that these events have similar formation mechanisms to the classical EBs found in active regions, with the reduced intensity enhancements found in the QS regions due to a weaker feeding magnetic flux. The observations also show that the quiet Sun Ellerman bombs (QSEBs) are formed at a higher height in the photosphere than the continuum level. Using simulations, we investigate the formation mechanism associated with the events and suggest that these events are driven by the interaction of magnetic field-lines in the upper photospheric regions. The results of the simulation are in agreement with observations when comparing the light-curves, and in most cases we found that the peak in the Ca II 8542 Å wing occurred before the peak in H α wing. Moreover, in some cases, the line profiles observed in Ca II are asymmetrical with a raised core profile. The source of heating in these events is shown by the MURaM simulations and is suggested to occur 430 km above the photosphere.

Key words: Sun: photosphere, Sun: chromosphere, Sun: magnetic fields, line: formation, line: profiles

1 INTRODUCTION

Ellerman Bombs (EBs) are prominent small-scale brightenings best observed in the far wings of H α . They were first reported by [Ellerman \(1917\)](#) as hydrogen bombs and were termed Ellerman bombs by [McMath et al. \(1960\)](#), while [Severny \(1956\)](#) termed them moustaches. They appear with a flame-like morphology, are 1000–2000 km in length and have vertical velocities of around 1 km s^{−1} with durations of 10 – 15 minutes ([Zachariadis et al. 1987](#); [Georgoulis et al. 2002](#)). EBs are generally observed near regions with relatively high concentrations of magnetic field, such as emerging flux regions and the penumbrae of sunspots ([Isobe et al. 2007](#); [Watanabe et al. 2008, 2011](#); [Rutten et al. 2013](#); [Vissers et al. 2013](#); [Nelson et al. 2015](#); [Reid et al. 2015, 2016](#), and references therein). Magnetic field configuration occurring in the photosphere dictates the morphology of the EBs ([Georgoulis et al. 2002](#); [Vissers et al. 2013](#); [Nelson et al. 2013](#); [Reid et al. 2015](#); [Tian et al. 2016](#); [Reid et al. 2016](#)). EBs are seen as enhanced intensities between 30% to 55% above average brightness in the wings of the H α line profile, often present above

the polarity inversion line ([Pariat et al. 2007](#); [Watanabe et al. 2008](#); [Reid et al. 2016](#)).

EBs are also observed in other lines. [Tian et al. \(2016\)](#) observed EB like events in Mn I 2795 Å, Mg II h and k lines, Ni II 1393.33 Å, and 1335.30 Å as enhancements in the wings rather than the core. EBs are observed in the Solar Dynamic Observatory's (SDO, [Pesnell et al. 2012](#)) Atmospheric Imaging Assembly (AIA, [Lemen et al. 2012](#)) 1700 channel as small brightenings. Out of the ten events that [Tian et al. \(2016\)](#) identified as UV bursts ([Peter et al. 2014](#); [Vissers et al. 2015](#)), seven were along the magnetic inversion line, and three were co-spatial with EBs. [Qiu et al. \(2000\)](#), show that there is a significant correlation with EBs in the H α wings at ± 1.3 Å and the UV continuum at 1600 Å. [Fang et al. \(2006\)](#) and [Pariat et al. \(2007\)](#), noted the presence of EBs in the Ca II 8542 Å lines. Spectropolarimetric observations done by [Pariat et al. \(2007\)](#), show that EBs are formed when opposite polarities merge giving rise to a cancellation of magnetic flux. As this cancellation occurs, plasma is heated and accelerated deep in the atmosphere and this is seen as a double-shaped hump in IRIS's Si IV, C II and Mg II lines. The total energies needed to produce EBs are estimated to be in the range of 10²⁷ to 10²⁸ ergs ([Georgoulis et al. 2002](#), and references therein), however in the IRIS observations the energy needed to

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drive the UV bursts is of the order of 10^{29} ergs (Peter et al. 2014). Rouppe van der Voort et al. (2016) studied Ellerman Bomb-like brightenings in the quiet Sun (QS) and suggested that these EBs can only be identified at the telescope's diffraction limit of $\lambda/D=0.14''$ at 6563 Å in SST data at a much lower intensity change, thus relaxing the 50% above average intensity requirement usually used to define EBs. Such QS observations of EBs were also reported by Nelson et al. (2017). Reid et al. (2017) report micro-flaring events that are in some cases similar to the classical definition of EBs and discussed the need for redefining EBs, based on signatures depending only on observations.

1.1 Magnetic concentrations and pseudo-EBs

Visser et al. (2013) classify an EB when the mean intensity enhancements are between 30% to 55% in the H α line wings, as compared to the average background line profiles. They further show that the bright grains, that are found simultaneously in the Ca II H, and the G-band images are bright network points. Such network bright points are driven by strong magnetic field concentrations (Sheeley 1969; Vrabec 1971; Harvey & Harvey 1973; Muller & Mena 1987; Hagenaar & Shine 2005).

Spruit (1976), suggest that these magnetic concentrations (MCs) are bright in the continuum of hot-wall radiation. Berger et al. (2004) and Rouppe van der Voort et al. (2005) indicate that the MCs rapidly evolve with complex morphologies. However, the MCs are found in the dark intergranular lanes and are only observed at a sub-arcsecond resolution (Title & Berger 1996). MCs are further observed in Mn I (Livingston & Wallace 1987), line wings of H α (Leenaarts et al. 2006) and the G-band (Leenaarts et al. 2006). They are less sharp in the Ca II H. In the DOT movies, Rutten et al. (2013) reports that MCs appear in the blue-wing of H α , suggesting down-flows. On comparing the signatures in H α and Na I D, they see MC shocks accompanied with blue-wing enhancements in H α .

Furthermore, Rutten et al. (2013) suggested that the mean intensity change in the wings of the H α line has to be at least 50% with respect to the average background line profile, and all EBs fainter than this should be considered as pseudo-EBs irrespective of the formation mechanisms. Such a definition suggests that the 3500+ EBs studied by Nelson et al. (2013) are pseudo-EBs. Additionally, Visser et al. (2015) suggested that the false-positives by Nelson et al. (2013), are because the regions studied were close to a decaying sunspot rather than an emerging sunspot. The most important difference between a MC and an EB is that EBs are related to reconnection.

1.2 Quiet Sun EBs (QSEBs)

QSEBs have a similar topology to EBs, such as a bright flame and lifetimes of a few minutes. Rouppe van der Voort et al. (2016) observed these events in the H α , Fe I 6173 Å, and Ca II 8542 Å, wavelengths in combination with IRIS and AIA/SDO. They found their intensities to be significantly lower than the active region EBs. However, they suggest that these features are also consequences of reconnection. Moreover, they also suggested that QSEBs are detected only when the data is of a high quality. Such data can be acquired from the Swedish 1-m Solar telescope (SST, Scharmer et al. (2003)) and is enhanced with the support of the adaptive optic system and image reconstruction techniques such as Multi-Object Multi-Frame Blind Deconvolution (MOMFBD, van Noort et al. (2005)). Rouppe van der Voort et al. (2016) identify 24 QSEBs

in a 4 Jul 2013 09:20 UT dataset and a further 21 QSEBs in a 4 Jul 2013 10:13 UT dataset. Furthermore, they describe QSEBs to have lengths between 150 to 360 km and widths ≈ 170 km. QSEBs are observed in positions $H\alpha \pm 1.3$ Å and last for a few minutes. QSEBs tend to have a predominantly bipolar topology, where after reconnection, both polarities seem to diminish. The intensity enhancement was below 40% in the H α line wing (in relation to the reference spectrum). Also, they concluded that these QSEBs are not observed in Ca II as they couldn't find significant evidence. In addition, Nelson et al. (2017), show the presence of QSEBs in their dataset.

1.3 Layout

We show certain cases of enhancements in the wings of H α , in the range of 10%–20% (above the QS average intensity) associated with the interaction of opposite polarities observed in Fe I. Most of these events are also observed in Ca II line wings with some events showing core enhancement. Such Ca II line wings enhancements were not reported with QSEBs before. We use light curves in H α , Ca II, Stokes-V and Stokes-I from Fe I, to investigate the evolution of QSEBs. Using a time series of magnetised photospheric models produced by the MURaM code (Vögler et al. 2005), we further analyse the character of plasma motions in intergranular magnetic field concentrations and in particular the formation height as seen in H α and Ca II 8542 Å. Such an approach provides a comprehensive understanding of the source of heating associated with these QSEBs.

The rest of the paper is organised as follows: Section 2 and Section 3, describe the methodology used in the project. Section 4 discusses the analysis with two detailed case studies. The simulation aspect is described in Section 5. The conclusion and key points from the paper are summarised in Section 6.

2 OBSERVATIONS

We investigate a quiet Sun disk center dataset, taken between 08:07:24 – 09:05:46 UT on 21 June 2012 using CRISP spectropolarimeter (CRISP, Scharmer et al. 2008) on the Swedish 1-m Solar Telescope (SST, Scharmer et al. 2003). We use QS observations in H α (6563 Å), Ca II (8542 Å), and Fe I (6302 Å). The FOV was centred in the QS at $[-3.1'', 69.9'']$. Fig 1 shows the location of the FOV against AIA 1700 Å channels. The subpanels show the zoomed-in view of SDO AIA 1600 Å and 1700 Å channels, with corresponding H α (6563 Å), Ca II (8542 Å) images. **Crosses ("X") represent the locations of 10 selected events and A corresponds to a unipolar event.** Multi-Object Multi-Frame Blind Deconvolution (MOMFBD) data reduction was performed using the method by van Noort et al. (2005). In H α , we observed at 10 line positions corresponding to ± 1.29 Å, ± 1.03 Å, ± 0.774 Å, ± 0.516 Å, ± 0.258 Å from the line center at 6563 Å (corresponding to Doppler velocities of ± 59 km s $^{-1}$, 47 km s $^{-1}$, 35 km s $^{-1}$, 23 km s $^{-1}$, 12 km s $^{-1}$. In Ca II (8542 Å) we observe at ± 0.495 Å, ± 0.440 Å, ± 0.384 Å, ± 0.330 Å, ± 0.275 Å, ± 0.219 Å, ± 0.165 Å, ± 0.110 Å, ± 0.054 Å with respect to the line center at 8542 Å. In Fe I (6302 Å) we obtained spectro-polarimetric observations only at one position at about -40 mÅ from the line core. The cadence of this dataset is 8 s.

3 DETECTION METHOD

We use the EB detection automation code "EBDATA" by Reid et al. (2016). The algorithm can detect and track magnetic concentrations. The algorithm relies on parameters of the magnetic concentrations

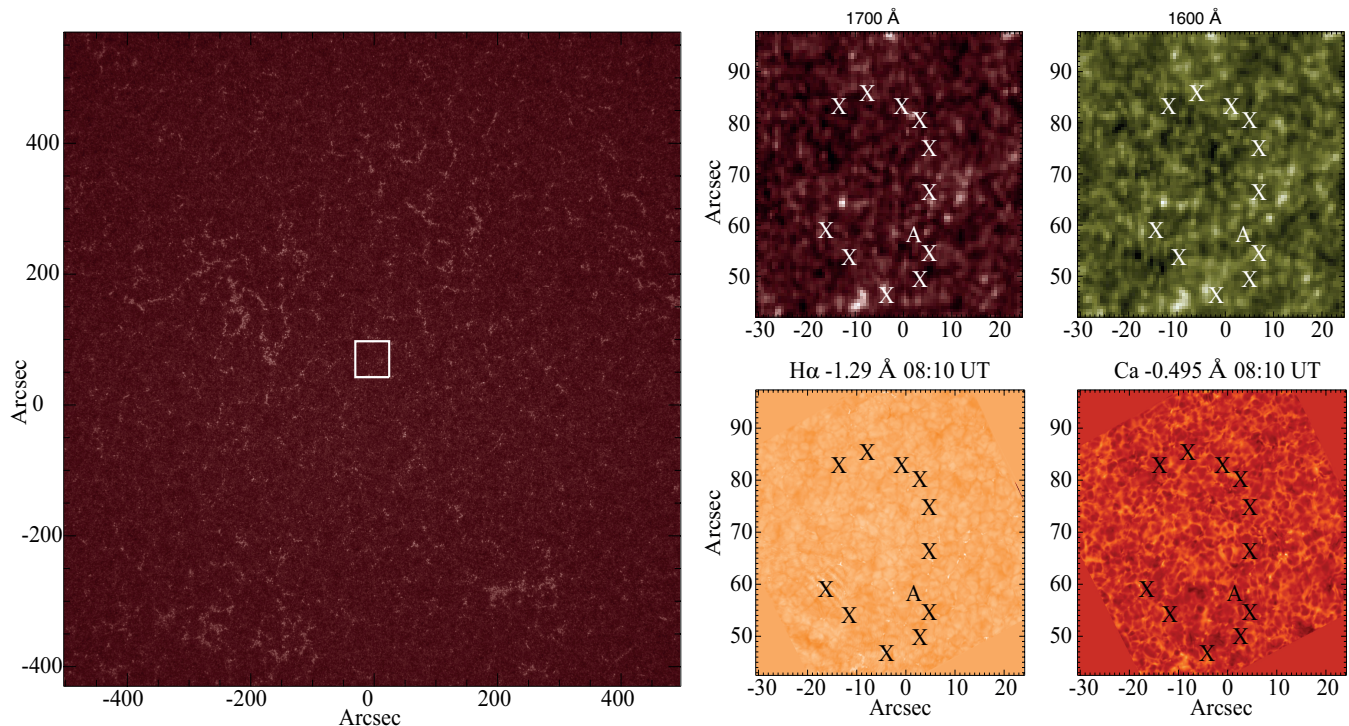


Figure 1. quiet Sun region observed using CRISP on SST. The left panel shows the location of the QS SST FOV overplotted against the AIA 1700 Å channel. The subpanels on the right panel, show the zoomed-in view of the SDO AIA 1600 Å and 1700 Å channels, with corresponding H α (6563 Å), Ca II (8542 Å) images obtained from SST. Crosses ("X") represent the locations of 10 selected events and A corresponds to a unipolar event.

such as intensity threshold, area, size and lifetime. The detection code also compares changes in intensity of the events with respect to surrounding intensity changes. In order to allow detections in the QS regions with low intensity changes (as compared to the background), we made minor changes to the algorithm. We define an event as a QSEB when the intensity contrasts, calculated in the H α wings, for at least one pixel is 10% more than the background average intensity. QSEBs evolve in time. This evolution is in terms of lateral motion as well as growth in size. The intensity of the grown area has to be greater than 10% that of the background average intensity, in both wings of the H α at ± 1.29 Å. The area of the QSEB, when it is fully grown has to be greater than 2 pixels (132 km). The line core in H α must remain unchanged (no more than 1% increase to account for variability as per Reid et al. (2016)). We detected 334 events, which are summarised in Fig 2. The left panel of Fig 2 shows a relation between the fractional change in the H α wing intensity and the apparent flux represented by the Stokes-V amplitude. The right hand panel shows a relation between the fractional change in the H α wing intensity and the rate of change of Stokes-V amplitude. Here, the maximum intensity is given as the maximum value in the wings of the detected pixels relative to the FOV average. The rate of change of Stokes-V amplitude is computed from change of Stokes-V intensities throughout the lifetime of the events. Most of the events identified by the routine were unipolar magnetic concentrations, shown by red circles and blue stars in the left panel of Fig 2. The rest were bipolar magnetic concentrations with possible EB-like wing enhancements. These are represented by black-crosses in the left panel of Fig 2. The events with less than 50 units of Stokes-V signal are termed as weak events and are represented by green squares in the top panel of Fig 2. In addition, the automated procedure also detected some

long lasting events with a strong unipolar region which would lie on the right hand side of the left panel of Fig 2 between "1000-1500". This unipolar event would correspond to a very low change in Stokes-V signal, and would lie near the "0" mark in the right hand panel. **The right hand panel of Fig 2, shows events which showed flux cancellation on the left hand side with negative flux signs. The events that showed emergence of flux are represented on the right-hand side of the plot, these show positive flux. All the selected events show flux cancellation.**

We manually selected 10 events from the detected events shown in Fig 2 satisfying properties of active region EBs, where magnetic flux cancellations are accompanied with wing enhancements in H α and Ca II 8542 Å. The events shown here are further selected by manual detection, which focused on 1.) interaction of QSEBs in Fe I 6302 Å Stokes-V wavelength, 2.) intensity enhancements in the H α wing positions and 3.) intensity enhancements in the Ca II 8542 Å wing positions. Also, we look for sudden intensity enhancements corresponding to the events. These intensity enhancements are smaller compared to regular active region EBs. After detecting events using the code, we manually checked whether they were formed above the interacting opposite polarity regions. Fig 2 shows scatter plots highlighting the comparison between properties of all detections and selected events. These events are labelled by the numbers 1–10 in Fig 2. For the selected events we present the snapshots of the H α wing position at -1.29 Å, Ca II 8542 Å at -0.495 Å, and Fe I 6302 Å Stokes-V as well as their appearance in the SDO AIA's 1600 and 1700 Å wavelengths in Fig 3. The white and black boxes are overplotted on the images. These boxes represent the region of interest, which are then used for further analysis. Three of the selected events have recurring intensity enhancements. such EB

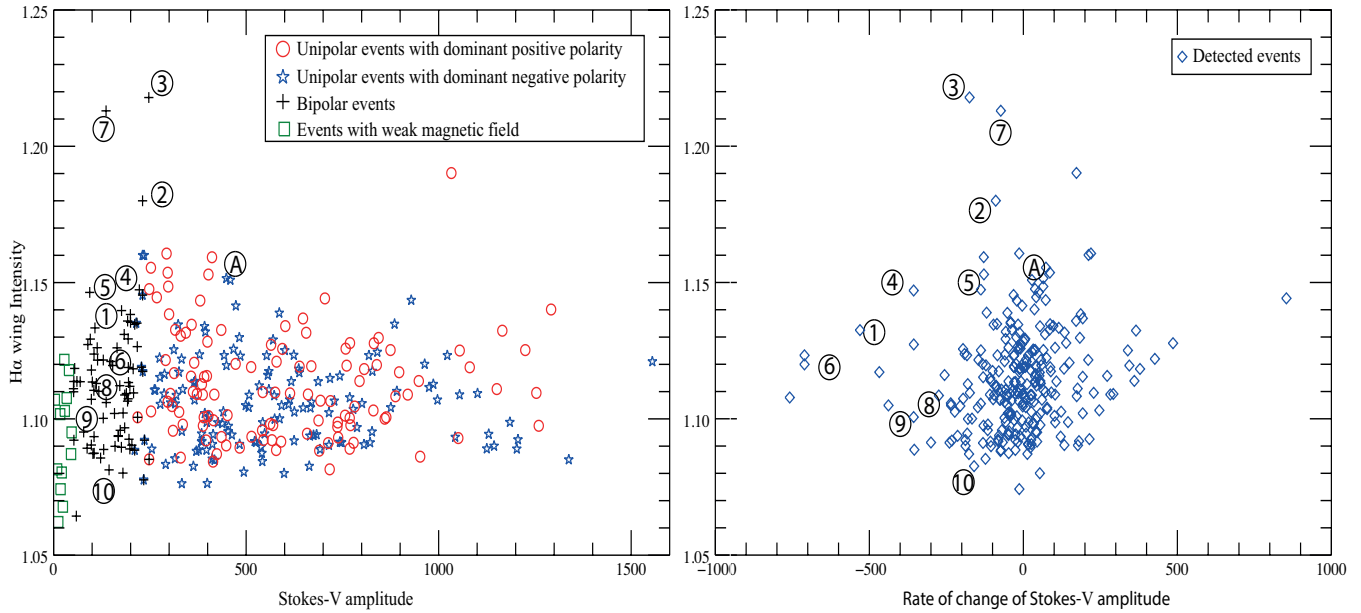


Figure 2. Scatterplot of detected events representing: in the left panel a relation between fractional change in the $H\alpha$ wing intensity and the Stokes-V amplitude, and in the right panel the relationship between the fractional change in the $H\alpha$ wing intensity and the rate of change of the Stokes-V signal. In the left panel the unipolar regions are represented by red-circles and blue-stars. Bipolar regions are represented by black-plusses. The events with less than 50 units of Stokes-V signal (weak events) are represented by green-squares. The manually selected regions are represented by numbers 1–10, and "A" represents a sample unipolar event. In the right panel, all the events are represented by blue diamonds with the selected events represented by 1–10 with a sample unipolar region represented as "A" (see text for details).

recurrence have also been seen in active regions (Qiu et al. 2000; Nelson et al. 2015; Reid et al. 2016).

4 ANALYSIS AND RESULTS

Fig 2 shows the manually selected events labelled 1–10. In the left hand side panel of Fig 2, the selected events lie in the region of bipolar detections with $H\alpha$ wing intensities between 1.06–1.23 and a high reduction in the Stokes-V intensity. A unipolar scenario event identified by the detection algorithm is labelled "A" in the plot. This event is not considered as a QSEB. This event shows maximum intensity similar to some bipolar detections however, since the event does not evolve during its lifetime, the rate of change of Stokes-V intensity is nearly zero. The same goes with other events identified by the detection code. The manually detected events are bipolar MCs with a significant reduction in the Stokes-V intensity. Thus the manually detected events stand out and have different properties from the other unipolar/bipolar events.

4.1 Light curves

In column 1 of Fig 4 and Fig 5, we plot $H\alpha$ (solid black) and Ca II 8542 Å (black dashed) light curves. The $H\alpha$ light curves are computed at ± 1.29 Å, and the Ca II 8542 Å light curves are computed at ± 0.495 Å. We plot light curves for the selected events by taking all pixels relating to the feature and summing the intensities in the blue and the red wing of the $H\alpha$ and Ca II 8542 Å lines. The

output is then divided by average value. The light curves show an impulsivity corresponding to EB-like line wing enhancements. We see enhancements in the $H\alpha$ and Ca II 8542 Å wings. The minimum intensities of the light curves correspond to times where the $H\alpha$ intensity contrast came back to the averaged background intensity. The maximum intensities correspond to our events observed in the $H\alpha$ and Ca II 8542 Å wing. We plot light curves in Stokes-V (solid black) and Stokes-I (black dashed) signal at -40 mÅ from the Fe 6302 Å line centre. The light curves are shown in column 2 (middle column). The light curves are plotted in orange and blue colours respectively in Fig 4 and Fig 5. We see that as one polarity disappears, the net-flux also reduces, indicating a cancellation of the magnetic field. In the column 3 of Fig 4 and Fig 5 we plot light curves corresponding to the SDO-AIA channels 1600 Å (solid black) and 1700 Å (black dashed). These are plotted by averaging all pixels corresponding to the events and dividing by the maximum intensity. These light curves represent the AIA 1600 and 1700 channels before and after the events. The green dotted vertical lines show the locations of the events as observed in CRISP.

From Fig 4 and Fig 5 we see that in the most cases (9 out of 10) the peak in intensity light curves occur in the wings of Ca II 8542 Å before $H\alpha$. The offsets at these positions are shown in Table 1. We can calculate the velocity of the down flow of the intensity enhancement. The approximate heights between the formation of $H\alpha$ and Ca II 8542 Å is assumed to be 500 km from models by Leenaarts et al. (2009) and Leenaarts et al. (2012). The propagation speed of the enhancements are approximately calculated using (distance/time) as 6 km s^{-1} . **For the EBs showing recurring activity**

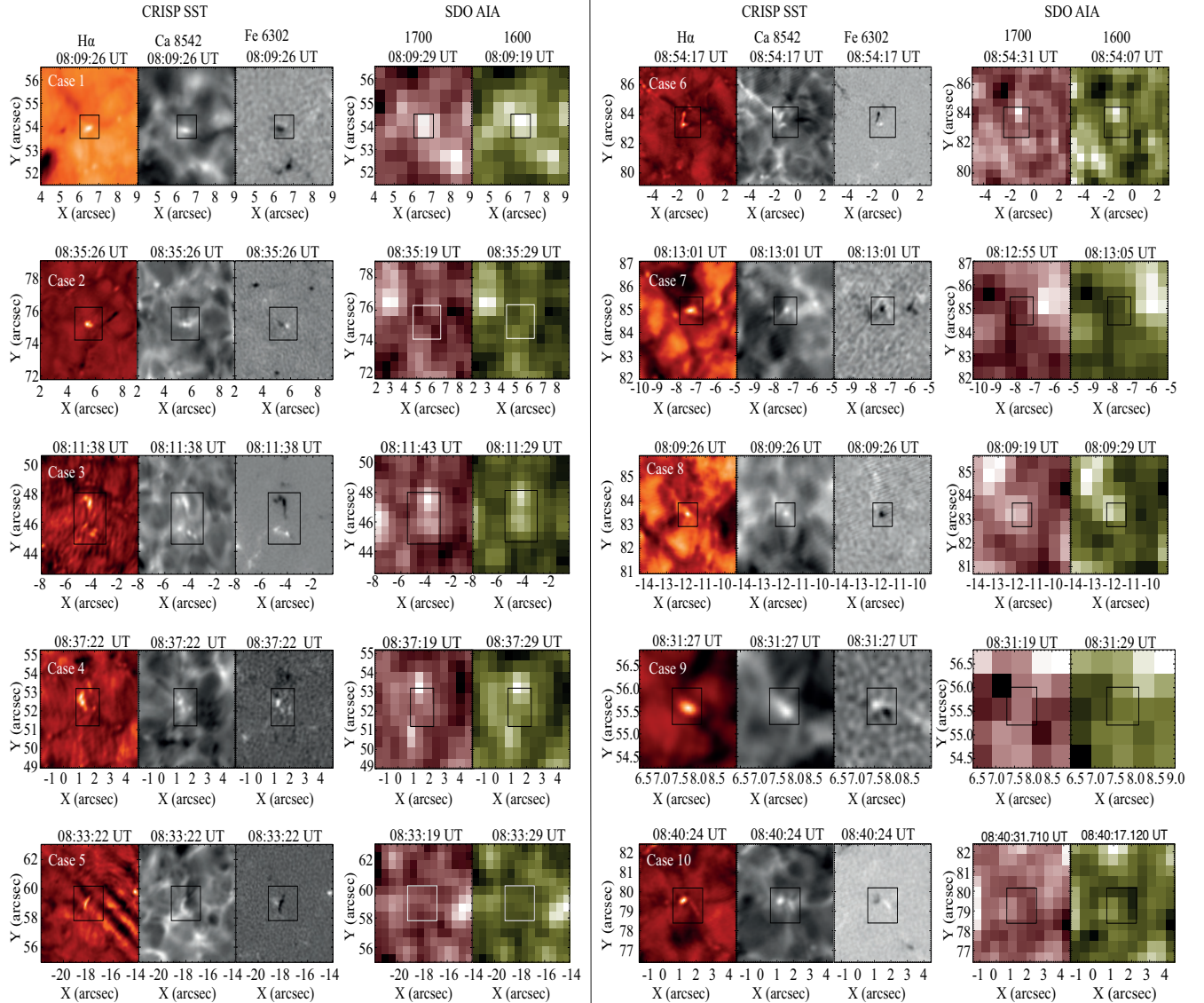


Figure 3. Snapshots of 10 possible QSEBs. The panels show the event in $H\alpha$ (first column) wing position -1.29 \AA , $\text{Ca II } 8542 \text{ \AA}$ wing position -0.495 \AA (second column) and $\text{Fe I } 6302 \text{ \AA}$ Stokes-V (third column). The AIA data from 1600 \AA and 1700 \AA is shown in the fourth and fifth columns respectively. The boxes overlapped on the images show the location of the QSEBs.

these velocities are calculated by using difference between most prominent peaks in the light curves of $H\alpha$ and $\text{Ca II } 8542 \text{ \AA}$. The impulsive nature of the events is associated with a corresponding decrease in the Stokes signals, thus suggesting that the intensity enhancements correspond to the magnetic flux cancellation possibility mentioned in Georgoulis et al. (2002). On comparing the 1600 and 1700 \AA light curves with the $H\alpha$ and $\text{Ca II } 8542 \text{ \AA}$ the intensity peaks observed in 1600 and 1700 \AA occur after the main intensity peaks observed in the $H\alpha$ and $\text{Ca II } 8542 \text{ \AA}$ (see cases 1, 3, 4 and 5). In some cases there is no apparent signature in the SDO channels (see case 6, 7 and 9). There are some cases which show a brightening in SDO-AIA channels at the location of the event. This brightening lasts for only one SDO frame (cases 2, 8 and 10). It is not clear whether these brightening corresponds to the event as it often appears after a delay. In addition, the light curves of cases 3, 8 and 10 in Fig 4 and Fig 5, show multiple impulsive bursts. Such behaviour is analogous to EBs observed near a large source

of magnetic flux. The light curves show that the QSEBs presented here form at an atmospheric level a few hundred kms above the photospheric continuum. This is shown by the different timings of the peak intensity between $\text{Ca II } 8542 \text{ \AA}$ and $H\alpha$ light curves.

4.2 Categories

By obtaining observations across multiple wavelengths, we can correlate the physics involved in events like EBs. In Fig 3, we show snapshots of 10 cases. The snapshots are taken at the $H\alpha$ wing position -1.29 \AA , $\text{Ca II } 8542 \text{ \AA}$ wing position -0.495 \AA , and $\text{Fe I } 6302 \text{ \AA}$ Stokes-V. The boxes represent the ROI. Based on the observations, events can be categorised into the following evolutionary characteristics: 1.) Single impulsive events involving reduction in Stokes-V after the intensity peaks in $H\alpha$ and $\text{Ca II } 8542 \text{ \AA}$ wings as in the light curves for cases 1, 2, 7 and 9. In case 9, the $\text{Ca II } 8542 \text{ \AA}$ and $\text{Fe I } 6302 \text{ \AA}$ Stokes-V track each other more closely than $H\alpha$, which

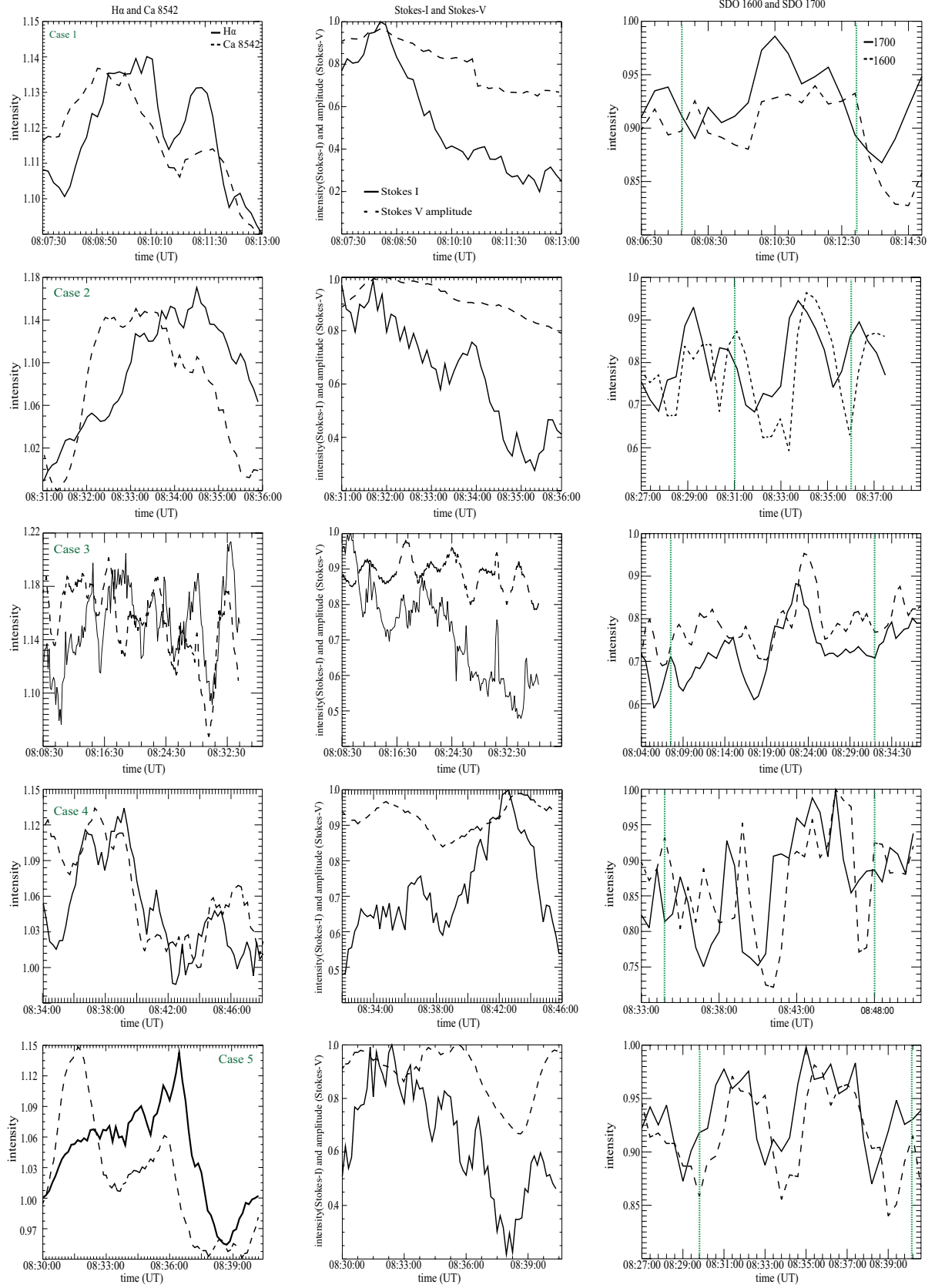
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Figure 4. Light curves for cases 1–5. Column 1: Light curves in $H\alpha$ (solid-black) wings at $\pm 1.29 \text{ \AA}$ and $\text{Ca II } 8542 \text{ \AA}$ (black dashed) wings at $\pm 0.495 \text{ \AA}$, Column 2: Light curves representing amplitude of in $\text{Fe I } 6302 \text{ \AA}$ Stokes-V (black dashed) and $\text{Fe I } 6302 \text{ \AA}$ Stokes-I (solid black). Column 3: Light curves in 1600 \AA (solid black) and 1700 \AA (black dashed) channels obtained from SDO-AIA. The vertical green dotted lines overplotted on the light curves obtained in the 1600 and 1700 channel represent the start time and end time of the event as observed in $H\alpha$ (solid-black) and Ca II .

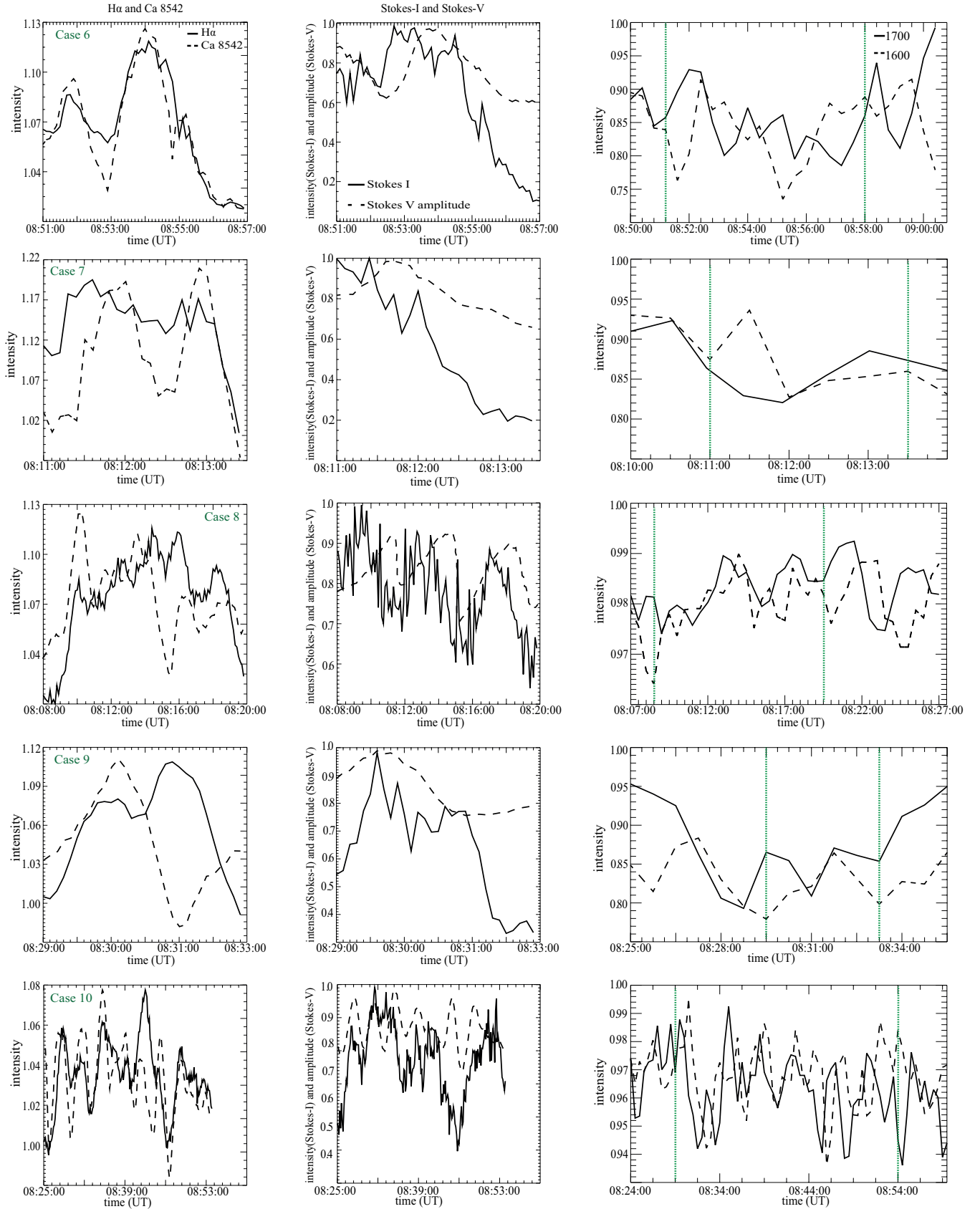


Figure 5. Light curves for cases 6–10. Column 1: Light curves in H α (solid-black) wings at ± 1.29 Å and Ca II 8542 Å (black dashed) wings at ± 0.495 Å, Column 2: Light curves representing amplitude of in Fe I 6302 Å Stokes-V (black dashed) and Fe I 6302 Å Stokes-I (solid black). Column 3: Light curves in 1600 Å (solid black) and 1700 Å (black dashed) channels obtained from SDO-AIA. The vertical green dotted lines overplotted on the light curves obtained in the 1600 and 1700 channel represent the start time and end time of the event as observed in H α (solid-black) and Ca II.

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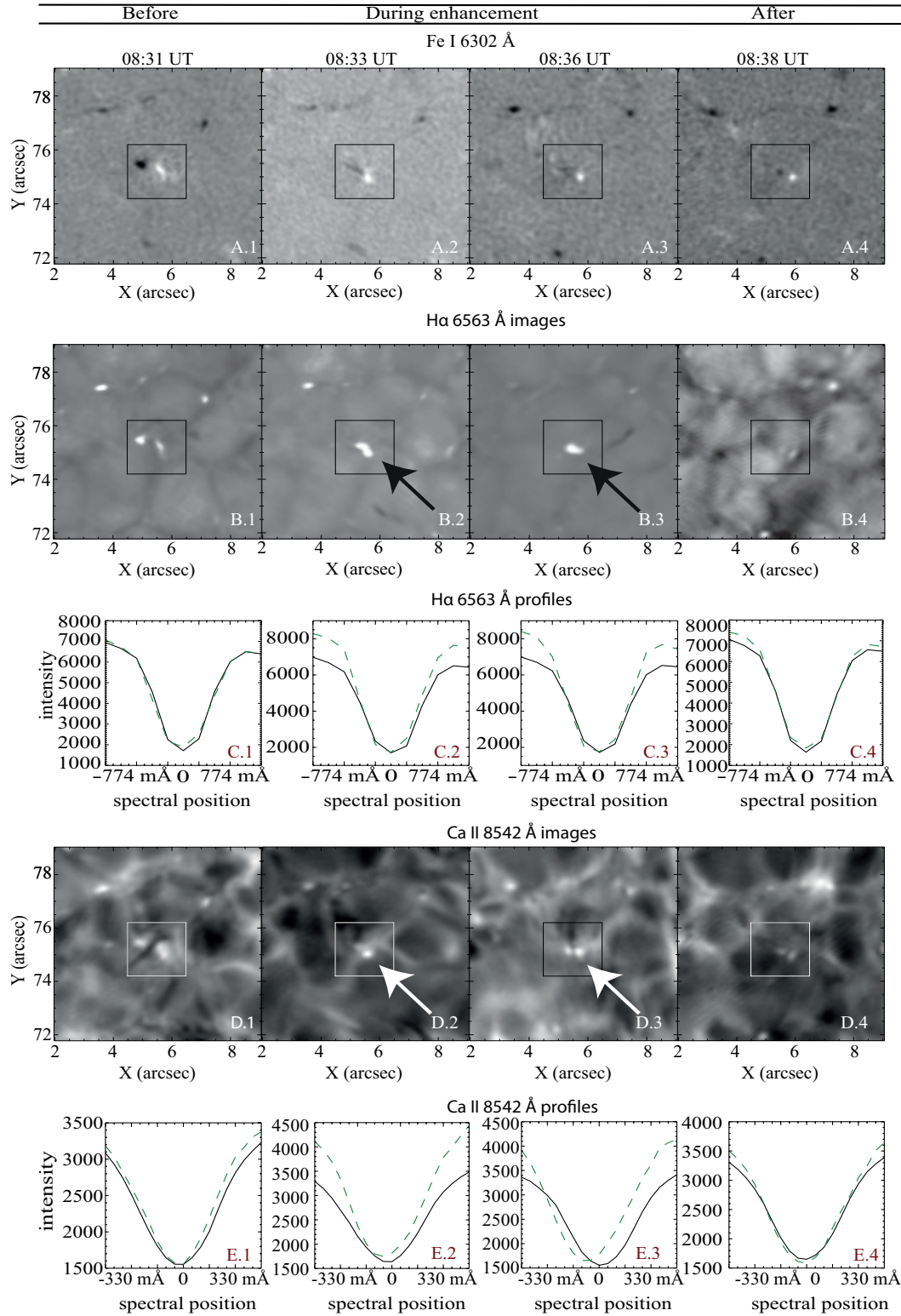


Figure 6. Evolution of the EB represented as Case 2. The top panel A1–A2, are snapshots of interacting opposite polarities as seen in Fe I 6302 Å. In the middle panel (B1–B2), we see an EB-like formation in the H α wings -1.29 Å and Ca II images at -0.495 Å (D1–D2). The locations of intensity enhancements are indicated by arrows. The line-profiles in H α are shown in panels C1–C4 and Ca II line-profiles are shown in E.1–E2. The green dashed lines represent the event and the averaged background line profile is represented by the solid black lines.

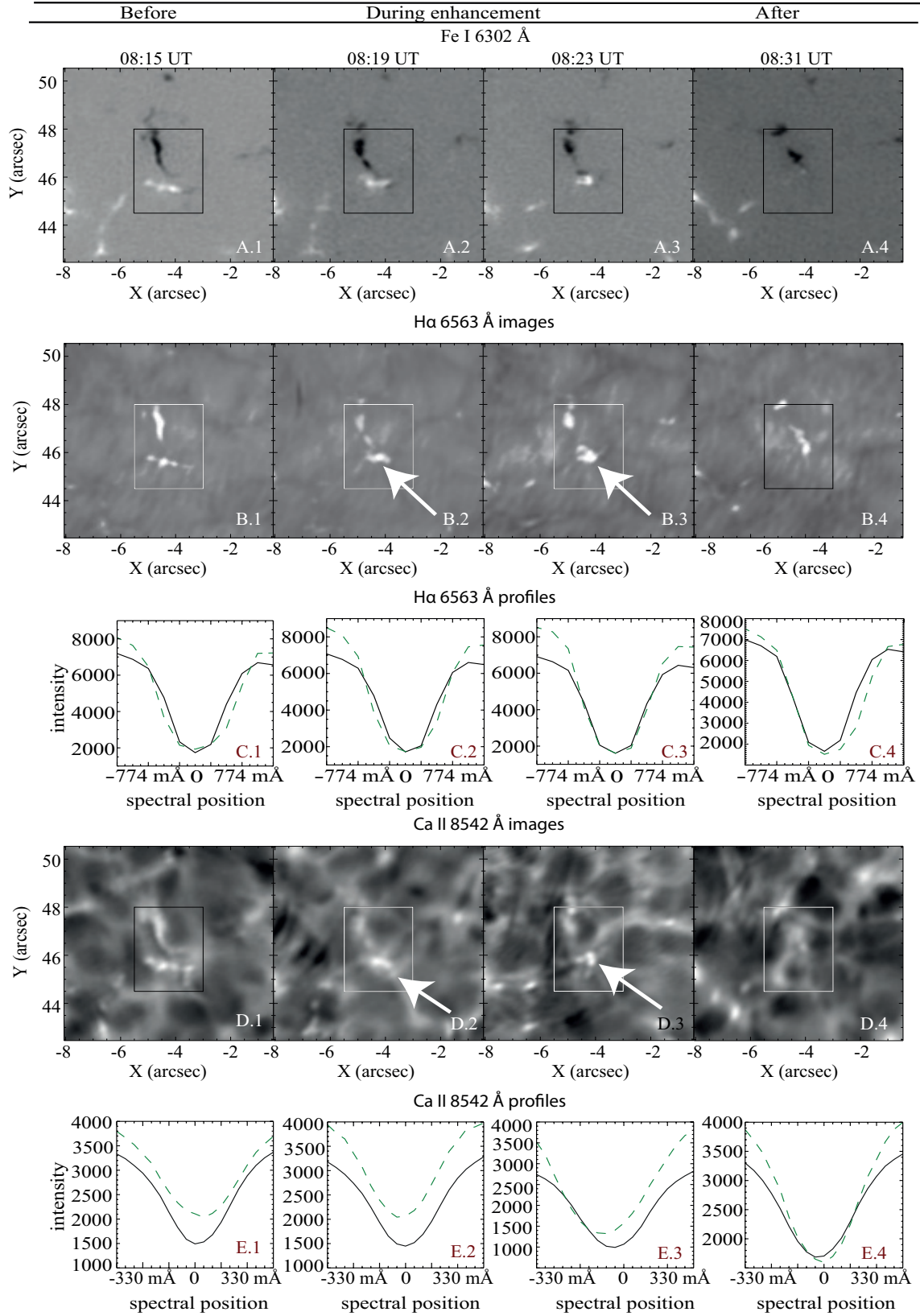


Figure 7. Evolution of the EB represented as Case 3. The top panel A1–A4 are snapshots of interacting opposite polarities as seen in Fe I 6302 Å. In the middle panel (B1–B4), we see an EB-like formation in the H α wings ~ 1.29 Å and Ca II images at ~ 0.495 Å (D1–D4). The locations of intensity enhancements are indicated by arrows. The line-profiles in H α are shown in panels C1–C4 and Ca II line-profiles are shown in E.1–E.2. The green dashed lines represent the event and the averaged background line profile is represented by the solid black lines.

Table 1. Calculation of plasma velocity.

Case No	ΔT s	Velocity of plasma. (Km/s)
1	60	8.33
2	120	4.15
3	420	1.21
4	80	6.25
5	260	1.92
6	10	50
7	-30	+16.4
8	300	1.51
9	640	0.78
10	330	1.51

is a common observational effect of dynamics related to magnetic concentrations. However, this reduction in intensity is followed by interaction of opposite polarity magnetic concentrations that gives rise to impulsivity. 2.) Events associated with reduction in Stokes-V signal where the two polarities keep on interacting (**cases 3, 8, and 10**). This reduction is observed with a repetitive impulsive nature in the $H\alpha$ and Ca II 8542 Å wings, during the time of the interactions. The presence of Ca II 8542 Å wing emissions in addition to $H\alpha$ emissions in all the cases suggest that such events are triggered in the lower chromosphere. The events studied show lifetimes of ~ 800 s with intensity change of $< 10\%$ in comparison to the average spectral lines. The velocity corresponding to the lateral motion of the selected events lies in the range of 0.3 km s^{-1} and 2.4 km s^{-1} . This velocity is computed by "EBDATA" detection algorithm. This velocity range matches with EBs found near active regions (Zachariadis et al. 1987).

4.3 QSEB morphology

In Fig 6, we show a small EB-like event. In the top-most row with panels A.1–A.4, we see that the two polarities interact continuously in the Fe I Stokes-V evolution. Such interaction gives rise to an enhanced emission in the $H\alpha$ wing images taken at ± 1.29 Å and Ca II 8542 Å line profile at ± 0.495 Å. The B.1–B.4 and D.1–D.4 of Fig 6 show snapshots taken in the $H\alpha$ wing position -1.29 Å and Ca II 8542 Å wing positions at -0.495 Å. We see typical EB topologies in both $H\alpha$ and Ca II 8542 Å images (see arrows in Fig 6). In sub-panels C.1–C.4 and E.1–E.4 of Fig 6, we show snapshots of the line-profiles with dashed green lines against background line profiles (the background line-profile is the average background across the FOV) shown in solid black lines. We see that there is a contrast change between $10\% - 20\%$ while comparing to the average $H\alpha$ spectrum. In Ca II 8542 Å we see that such events have higher contrast changes from 20% to 40% as compared to the average Ca II 8542 Å spectrum. However, the line profile in Ca II 8542 Å is asymmetric, see panels E.3 of Fig 6.

In Fig 7, we show an event which involves two magnetic concentrations interacting for ~ 15 mins. The photospheric flux cancellation is followed by repetitive emissions in the $H\alpha$ and Ca II wings. In the subpanels A.1–A.4, we see opposite polarities interacting in the Fe I 6302 Å line core images, in the evolution. The interaction between negative and positive polarity causes the weaker polarity to be annihilated over the evolution (not shown here). Furthermore, the subpanels A.1–A.4, show evolution of the two polarities where the positive polarity is seen to diminish in size in A.4 as compared to A.1. This merging and interaction gives rise to multiple intensity

peaks, seen in the $H\alpha$ wing position -1.29 Å and Ca II 8542 Å -0.495 Å images (see subpanels B.1–B.4 and D.1–D.4). Below both the $H\alpha$ and Ca II 8542 Å images, we show snapshots representing line-profiles with dashed green lines against solid black lines which represent the average background spectrum (see subpanels C.1–C.4 and E.1–E.4).

4.4 A sample unipolar event

The EBDATA algorithm detected 334 events, out of which 10 were selected for detailed analysis. We discuss here the evolution of a unipolar event that was discarded as a "false positive". The unipolar event is labeled "A" in the Fig 2. The snapshots of the evolution of this unipolar event are shown in subpanels A.1–A.3 of Fig 8 with larger negative polarity in size (sub panels A.1–A.3 of Fig 8). In subpanels B.1–B.3, we show a series of the $H\alpha$ images. We see enhancements in the intensity at locations where the unipolar flux region combines. Such intensity enhancements are also seen in subpanels D.1–D.4. In both the $H\alpha$ and Ca II 8542 Å images, we see EB-like wing enhancements. The $H\alpha$ and Ca II 8542 Å line profiles show a similar behaviour of emission as compared to previous examples (see the snapshots of line profiles seen in subpanels C.1–C.3 and E.1–E.3 respectively). However, Ca II 8542 Å line profile shows a strong blue-shifted line profile with core enhancements. Such events could be due to shearing reconnection, low-resolution imaging fails to spot the opposite polarity, or they could be driven by braided reconnection. Furthermore, Fig 9 shows no clear relation between the $H\alpha$ (solid black) and Ca II 8542 Å (black dashed) light curves. The examples here show impulsivity observed in $H\alpha$, which may or may not be related to the QSEBs. Hence we have ignored such detections.

The observational diagnosis indicates that in disk-centre viewing along the radial direction, only the top of an EB is seen, which shields what lies underneath as noted in simulations by Danilovic (2017). Thus there is an absence of flame-like topology here. Also, in comparison with the 1600 and 1700 Å channels, we see no particular correspondence with the EB signatures observed in $H\alpha$ and Ca II 8542 Å. However, we do note that in some cases there is some brightening that occurs after the initial EB brightening that could be related to these events. **A possible explanation for the lack of UV enhancement could be due to the lower spatial resolution of the SDO AIA instrument, or the lower height at which the UV continua form.** It is analogous to reports by Rouppe van der Voort et al. (2016). Danilovic (2017) conclude that the strongest brightening corresponds to a significant temperature and density increase that occurs at the site of the cancellation of two magnetic features of opposite polarities. Furthermore, the authors also highlight that unipolar regions are also strong EB candidates when accompanied by flux cancellation. This highlights that many detected unipolar regions could be an EB candidate. Georgoulis et al. (2002) suggest that flux cancellation is possible in unipolar regions by shearing reconnection. Furthermore, Hansteen et al. (2017) using BiFROST simulations suggest a weak brightening in Si IV associated with EBs.

5 MURAM SIMULATIONS

The aim of this simulation is to understand the formation mechanisms related to QSEBs. They complement the observations as the

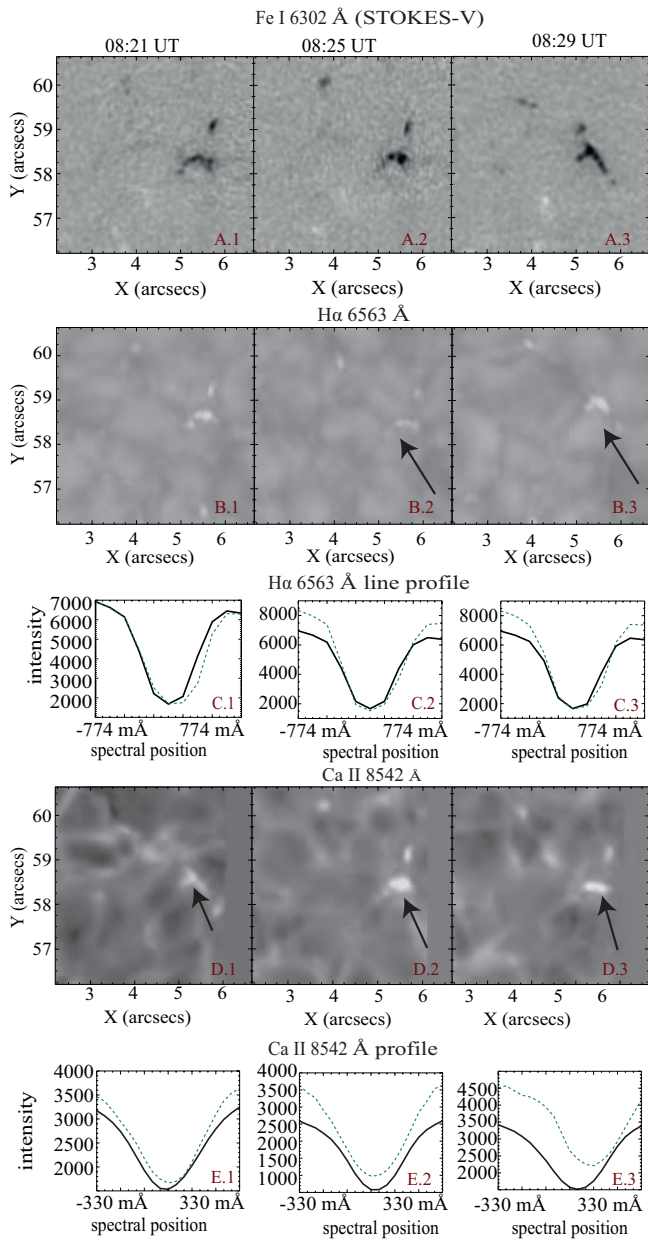


Figure 8. Top panels shows a unipolar event with dominant negative polarity (indicated by "A" in Fig 2). The top panel A1–A2, represents snapshots of interacting opposite polarities as seen in Fe I 6302 Å. In the middle panel (B1–B2), we see a EB-like formation in the H α wings ± 1.29 Å and Ca II 8542 Å images at ± 0.495 Å (D1–D2). The locations of intensity enhancements are indicated by arrows. The line-profiles in H α are shown in panels C1–C3 and Ca II line-profiles are shown in E.1–E.2. The green dashed lines represent the event and the averaged background line profile is represented by the solid black lines.

simulations performed in MURaM (Vögler et al. 2005) tell us where these events are formed. The MURaM code is used to perform simulations of the interaction of the magnetic field concentrations in the solar photosphere. This particular setup is for a QS region. The numerical setup for these set of observations is similar to the one described in Nelson et al. (2013). The spatial resolution of the box is $25 \text{ km} \times 14 \text{ km} \times 25 \text{ km}$. The temporal resolution of the simulation is 50s. A positive-negative "checkerboard" vertically-directed

Signatures of quiet Sun reconnection events. 11

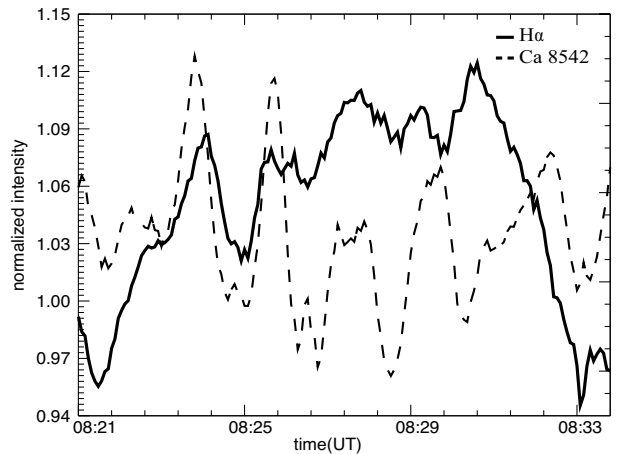


Figure 9. Light curves for unipolar region shown in Fig 8. Light curves in H α (solid-black) wings at ± 1.29 Å and Ca II 8542 Å (black dashed) wings at ± 0.495 Å.

magnetic field, with the unsigned strength of 200 G is added to a well-developed non-magnetic photospheric convection snapshot. Then the computational domain is set to evolve for a small number (2-5) of granular lifetimes. During the evolutionary period, most of the magnetic field cancels out, leaving some substantial magnetic field concentrations of opposite polarities in the intergranular lanes of the simulated photospheric granulation. These magnetic field concentrations move along the intergranular lanes occasionally coming in proximity to each other and reconnecting.

Fig 10, shows one such evolution for magnetic field concentrations at the approximate height of the photosphere, and is represented by subpanels A1–A4. Here we see two magnetic concentrations of opposite polarities interacting with each other. The time-stamps are separated by 50 s. Approximately 150s into the simulation, one of the polarities cancels out. This is similar to what we observe in Fe I 6302 Å for all events (see sub-panels A1–A4 of Fig 6 and Fig 7). The corresponding magnetic field cancellation rate at the photosphere and at the level 430 km above the photosphere is shown in Fig 11. Here, the left panel corresponds to the subpanels A1–A4 of Fig 10. These magnetic flux curves are plotted by summing the magnetic flux in the opposite polarities, as the polarities evolve in time.

Furthermore, the intergranular magnetic field concentrations expand into the higher layers of the simulated solar atmosphere due to a magnetic-thermal pressure balance and thermal pressure decrease with height. Such evolution of intergranular magnetic field concentrations are shown in subpanels B1–B4 of Fig 10. These timestamps are taken at 430 km above the photosphere and are separated by 50 s. The right panel of Fig 11 shows a magnetic flux cancellation rate at this level.

Due to the geometry of the magnetic field in the simulations, it is expected that the reconnection process evolves in time from the top of the simulation domain towards the solar interior, with the reconnection point moving downwards. This instant of reconnection is seen as an enhancement in the simulations. Fig 12, shows temperature maps taken at the continuum level and lower chromosphere/upper photosphere (~ 430 km above the photosphere). Each subpanel is separated by 50s. The red arrows show the locations which indicate heating (dark/black colour). Here we see that the temperature rise occurs in subpanel B1 and continues throughout the evolution. However, in the subpanels representing the photo-

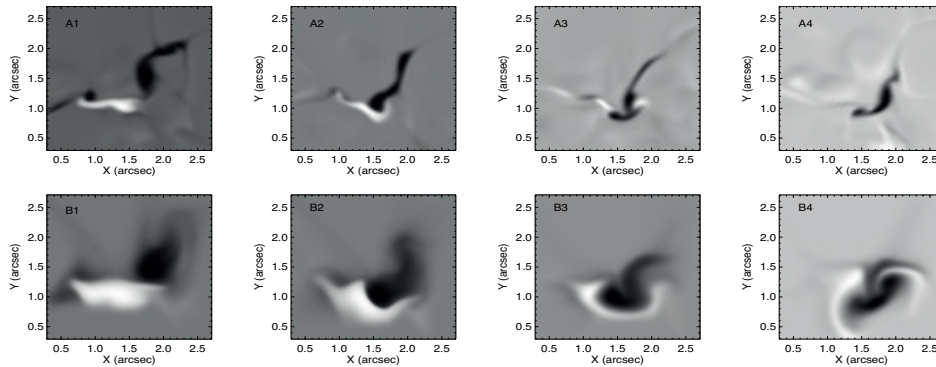


Figure 10. Evolution of magnetic field concentrations in the MURaM simulations. The time-stamps are separated by 50s. The top row with subpanels A1–A4 shows interacting opposite polarity magnetic concentrations at the photospheric levels. The bottom row with subpanels B1–B4, show the responses of the interactions to magnetic concentrations at ~430 km above the photosphere.

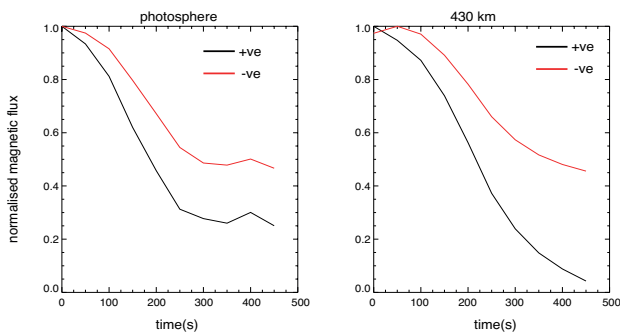


Figure 11. The rate of change in magnetic flux at Fe I 6302 Å as simulated in the MURaM simulations for the regions represented in Fig 10. The left panel shows the rate of change in magnetic flux at the photospheric continuum level and the right panel shows rate of change in magnetic flux at 430 km above the photosphere.

sphere, the enhancement here is relatively small and appears ~100 s later. Such behaviour matches with the observations. The temperature curves are plotted in Fig 13, the solid red line corresponds to the photospheric continuum level while the solid black line corresponds to the lower chromospheric (upper photospheric) level. The temperature peaks at ~250 s at 430 km above photosphere and at ~350 s for the photosphere.

The time for the reconnection point to move downwards can be estimated by a similar method described in [Keys et al. \(2013\)](#). Here the authors calculate a velocity of 1.8 km s^{-1} for bright point motions. The mean horizontal speed in the photosphere in the reconnection region, as the simulations show, is $\sim 4 \text{ km s}^{-1}$ ($2 \times 1.8 \text{ km s}^{-1}$). For the sake of simplicity, it can be assumed that the reconnecting magnetic field concentrations move towards each other with this speed. The expansion factor of the intergranular magnetic flux tubes is about 2 between the continuum formation layer and the layer 600 km above it, and the magnetic field concentration size at the photospheric layer is about 200 km. Assuming straight field lines, the fields from two flux concentrations touch at +600 km when there is a 2 km gap between them at the continuum formation layer. It takes about 50 s for the flux concentrations to touch at the continuum formation layer during which time the contact point above has moved down by 600 km. Therefore, the time for the reconnection point to move, from the +600 km layer down to the continuum formation layer is also about 50 s. The same calculations work if it takes 100 s to move from the upper photosphere to the continuum

layer. The calculation depends on the relative motion speed of the opposite polarity magnetic field concentrations given in Fig 13. A simple model for this calculation is sketched in the diagram in Fig 14.

This is a simplistic calculation, and it does not represent the whole complexity of the dynamic process occurring in the simulation domain. However, it gives us a clue of the scale of the reconnection timescales, and consequently the delay between the upper photospheric and lower chromospheric signals in the Stokes-I profiles of the corresponding absorption lines (e.g. H α 6563 Å and Ca II 8542 Å).

6 DISCUSSIONS AND SUMMARY

We use observations from H α , Ca II 8542 Å, and Fe I 6302 Å spectral lines, to investigate the lower solar atmosphere. In Fe I 6302 Å (~40 mÅ), we observe the Stokes-V signal, which is similar to a magnetogram. We observed several events that gave rise to impulsive, flare-like enhancements in the wings of H α and Ca II 8542 Å. These events are associated with the interaction of opposite polarities in Stokes-V of Fe I 6302 Å. The aim of the paper is to show that QSEBs are observed as low intensity contrast events. Fig 2, shows our selected events correspond to a reduction in the Stokes-V intensity accompanied by a maximum H α wing intensity. The peak in the H α wing intensity is $\approx 20\%$ above the average background or less. When compared to the other events in the detection algorithm these events stand out. Thus we have presented QSEBs with less than 20% intensity increase that satisfies various cancellation models discussed by [Georgoulis et al. \(2002\)](#) and have all the signatures of EBs found in an active region. However, they show a low-intensity impulsive nature (see sub-panels C1–C4 of Fig 6 and Fig 7). We have used MURaM simulations to understand these events. The sudden enhancement in the wings of the H α line and Ca II line profiles suggest a physical nature similar to that of EBs. We propose that the reason for the low-intensity contrast of QSEB compared to active region EBs is due to the weaker flux cancellation and the subsequent energy transferred to radiative energy is lower than in regular active region EBs. EBs present near a sunspot have a characteristic recurrent flame-like emission, which recurs with simultaneous H α , and Ca II 8542 Å wing enhancements. We see such recurring emissions in QSEBs that have a well-defined EB-like morphology (see Fig 7). In addition to the H α signatures, we see an increase in both the core and the wings of the Ca II line profile. Ca II

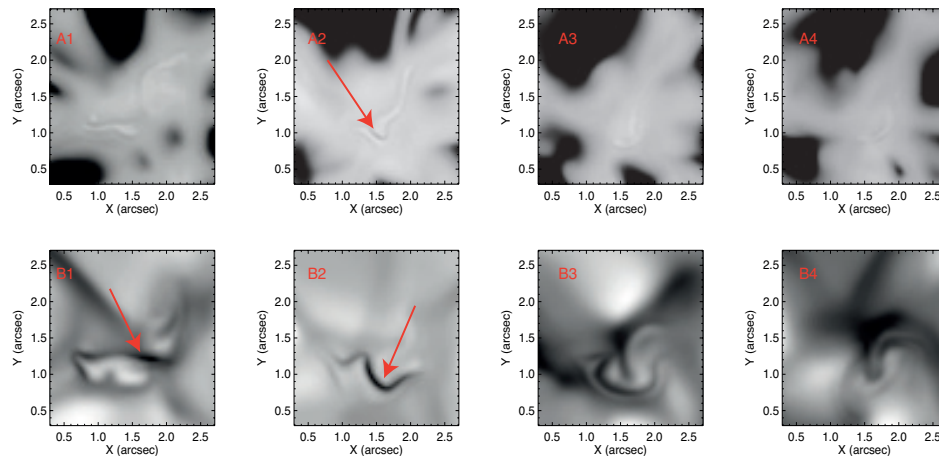


Figure 12. Evolution of temperature, corresponding to panels in Fig 10 from the MURaM simulations. The time-stamps are separated by 50s. The top row with subpanels A1–A4, shows temperature change at the photospheric level. The bottom row with subpanels B1–B2, shows the temperature change at ~430 km above the photosphere.

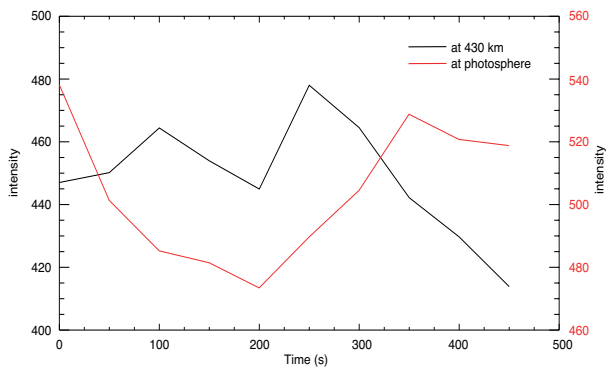


Figure 13. Evolution of intensity, corresponding to panels in Fig 10 from the MURaM simulations. The timestamps are separated by 50s. The top row with the subpanels A1–A4, shows temperature change at photospheric levels. The bottom row with subpanels B1–B2, shows the temperature change at 430 km above the photosphere.

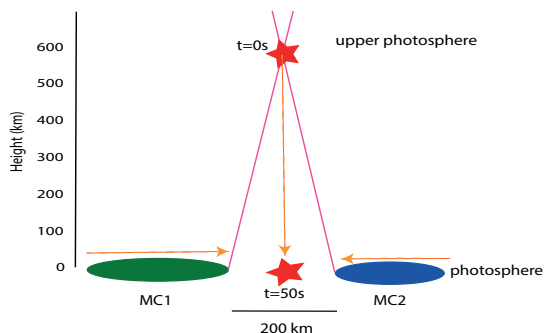


Figure 14. Simple cartoon explaining the observed and simulated scenario. Here, we see two MCs separated by 200 km. The arrows indicate the direction of apparent motion. The pink solid lines show the possible straight magnetic field lines. The stars indicate two flaring regions. The flaring at 600 km (upper photosphere) occurs 50s before the flaring at the continuum level. The image is not scaled.

profiles associated with the QSEBs are also asymmetrical compared to the $H\alpha$ profile.

Another aspect of our observations is the presence of a temperature increase corresponding to the QSEBs. This temperature increase is especially seen in $\text{Ca II } 8542 \text{ \AA}$ wing emissions. The light-curves (see Fig 4 and Fig 5), show that in most of the cases we see $\text{Ca II } 8542 \text{ \AA}$ wing emissions occurring before the $H\alpha$ wing emissions. The temperature increase further indicates that the increase in emission intensity occurs higher in the upper photosphere and the effects propagate downwards. Such morphology is also observed in the MURaM simulations (see Fig 12 and Fig 13). Here we see that the temperature increase occurs higher in the atmospheric layer (at 7000K or 430km above the photospheric continuum) and occurs before a temperature rise in the photospheric continuum level. This model is further supported by the fact that the SDO channels formed at the continuum level show intensity peaks after the intensity peaks observed in the chromospheric lines of $H\alpha$ and $\text{Ca II } 8542 \text{ \AA}$.

Furthermore, our simulations indicates that only a small temperature increase in the lower photosphere is required to reproduce the observed line profiles. This temperature change occurs at the continuum layer 480 km above the assumed photosphere (see Fig 13). These simulation gives us a clue of the scale of the reconnection timescales, and consequently the delay between the upper photospheric and lower chromospheric signals in the Stokes-I profiles of the corresponding absorption lines (e.g. $H\alpha$ 6563 \AA and $\text{Ca II } 8542 \text{ \AA}$).

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